## Astro 210 Lecture 17 October 4, 2010

#### **Announcements**

- ★ Guest Instructor today: Prof. Athol Kemball a real observer! and expert on radio astronomy!
- HW5 available, due next Friday
- required Night Observing this week check online for schedule and weather info download & bring question sheet

Last time: gasses under a microscope gasses: collections of particles, e.g., atoms, molecules

www: gas law simulation

- in space: "elbow room" empty space between particles
- in time: constant random motion collisions exchange energy, momentum
- individual particle velocities random, changing but distribution (=historgram, bar graph) of particle speeds is set by temperature
  - $\rightarrow$  average speed  $v_{\rm rms} = \sqrt{3kT/m}$

### Microsopic View of a Piston

Now consider a large number of gas particles

- in a sealed volume
- with a *piston* of area A

from microscopic viewpoint: piston constantly bombarded by gas particles if let free—would be pushed away

to resist bombardment, must push on piston = exert force F define pressure

$$P = \frac{\text{force on piston}}{\text{area of piston}} = \frac{F}{A} \tag{1}$$

Q: how to intensify bombardment = pressure on piston? Hint-more than one way to do this

#### **Pressure**

collisions with walls  $\rightarrow$  momentum transfer  $\rightarrow$  force  $\rightarrow$  pressure www: piston simulation

#### ideal gas

pressure P, volume V, total number N of particles and absolute (Kelvin) temperature T all related by ideal gas equation of state:

$$PV = NkT (2)$$

- N counts individual particles, typically very large! alternatively: can count in units of moles of particles i.e., in units of  $N_{\text{Avo}} = 6 \times 10^{23}$  then # moles is  $n_{\text{moles}} = N/N_{\text{Avo}}$  and  $PV = n_{\text{moles}}RT$ , where  $R = N_{\text{Avo}}k = 8.3$  Joules mole<sup>-1</sup> K<sup>-1</sup>
- since density  $\rho = M/V = mN/V$ , with m gas particle mass can rewrite  $P = \rho kT/m$

## **Planetary Atmospheres**

#### Terrestrial Atmospheres:

- atmospheres are tiny (or zero!) fraction of planet mass
- no light gasses (H<sub>2</sub>, He),
   only heavier N<sub>2</sub>, O<sub>2</sub>, CO<sub>2</sub> (if anything!)

#### Jovian Atmospheres:

- a significant fraction of planet mass
- mostly H<sub>2</sub>, He; some heavier species

Q: why the difference? what factors important?

#### competition: gravity versus thermal motion

gravity  $\rightarrow$  keep particles thermal motion  $\rightarrow$  run away

• gravity  $\rightarrow$  ecape speed  $v_{\rm esc} = \sqrt{2GM/R} = 2.4 \times 10^3 \ {\rm m/s} \ {\rm for \ Moon}$ 

• thermal motion  $\rightarrow$  avg thermal speed  $v_{\rm rms} = 2.6 \times 10^3$  m/s for H on Moon  $v_{\rm rms} > v_{\rm esc}$ : many atoms can escape H lost from Moon (check!)

really: not all particles have same speed

so always some escape

real question: timescale

to keep atm for age of SS

need  $v_{\rm esc} \ge 6v_{\rm rms}$  (rule of thumb)

species	$v_{ m esc}/v_{ m rms}$	lifetime
H on moon	0.9	hours
H on earth	2.1	days
sodium on moon	4	$\sim 10^3$ yrs
O <sub>2</sub> on earth	12	$\to \infty$ : no escape!

www: Sodium lunar ''atmosphere'' solar UV photon on Moon ''soil''  $\to$  thin vapor of sodium ejected but replenished

# LET'S GO: SOLAR SYSTEM

#### The Earth

astro-trivia: Earth logo/icon is  $\oplus$  recall:  $R_{\oplus} = 6.4 \times 10^4$  m = 6400 km get mass from  $g = GM_{\oplus}/R_{\oplus}^2$  (need G!)  $\to M_{\oplus} = 6.0 \times 10^{24}$  kg

#### Average density:

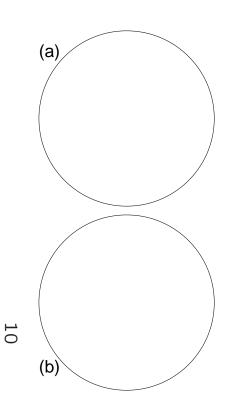
 $\rho_{\rm av, \oplus} = 3/4\pi M_{\oplus}/R_{\oplus}^3 = 5,500 {\rm kg/m^3}$  between rocks and iron  $\to$  some of both

Orbit around Sun: ellipse, eccentricity  $e_{\oplus} = 0.017$ 

#### iClicker Poll: Earth's Orbit

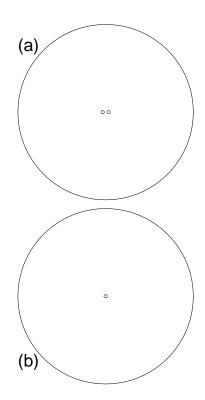
Vote your conscience!

Below: one ellipse with  $e=e_{\oplus}=0.017$ , one circle (e=0) Which of these has the eccentricity of Earth's ellipse?



## **Survey Says**

## Foci labelled



## **Seasons and Eccencricity**

Due to elliptical orbit, Earth-Sun distance changes

$$r_{\sf ap} = (1 + e)a_{\oplus} = 1.017 \; {\sf AU}$$

$$r_{\rm peri} = (1 - e)a_{\oplus} = 0.983 \text{ AU}$$

...but not by much! 3.2% swing!

if this were the whole story:

temperature swing by  $\approx 2\% \times 300 \text{ K} = 6^{\circ} \text{ C} = 11^{\circ} \text{ F!}$ 

also: if yearly temperature variations were due to eccentricity...

Q: how should the seasons compare in the N and S hemispheres?

Q: what is actually observed?

Q: what does this imply about season origin?

## Origin of the Seasons

seasons are **exactly opposite** in N and S hemispheres i.e., right now is spring in S. America, will soon be summer

so: season origin must distinguish hemispheres

recall: Earth's spin axis tilted w.r.t. orbit by 23.5°

- when one hemisphere tilted closest, the other farthest: this is summer/winter
- ullet when tilt ot Earth-Sun radius: fall/spring

Note: Earth closest to Sun in northern winter!

→ eccentricity has tiny effect on temperature

#### Earth's Interior

crust: 16-40 km

mantle:  $\sim$  3000 km

outer core:  $\sim$  2,200 km

inner core:  $\sim 1,200$  km

inner core: solid. Fe, Ni

outer core: liquid. Fe, Ni

mantle: "plastic". Fe, Mg, Si, O

crust: solid. ocean basins—basalt: O, Si, Al, Mg

continental plates-granite: O, Si, Al, Na, K

heaviet elements lowest → settling ("differentiation")

Q: how do we know?

How do we know?

"Refraction" of Earthquakes

Demo: slinky

Earthquakes "emit" waves  $\rightarrow$  use seismographic info to learn earth structure diagram

#### **Plate Tectonics**

www: plot of earthquake sites

www: plot of volcano sites

www: plate locations

crust not a sigle rigid solid but collection of "plates" motions in mantle (convection) cause plate motion sketch convection currents

www: satellite laser ranging

www: VLBI: radio telescopes used to detect motion

www: drift animation

plates move, and interact: collide, slide, buckle

⇒ "plate tectonics"

leads to observed geological features

www: Mountian

www: volcano (Kilauea, HI)

www: San Andreas

Earth is evolving!