Astro 501: Radiative Processes
Lecture 34
April 19, 2013

Announcements:
- **Problem Set 10** due 5pm today
- **Problem Set 11** last one! due Monday April 29

Last time: absorption line formation
Q: at high resolution, what qualitative and quantitative information does a line give?
Q: what must be known to extract this information?

line *transmission profile* $F_\nu / F_\nu(0) = e^{-\tau_\nu}$
directly measures optical depth $\tau_\nu \approx \sigma_{\ell u} N_\ell$

So *if we assume we know the spectral shape* $F_\nu(0)$
of the background source across the line transmission profile
then measure product $\sigma_{\ell u} \tau_\nu$

but the absorption cross section is

$$\sigma_{\ell u}(\nu) = \pi e^2 / m_e c \; f_{\ell u} \; \phi_{\ell u}(\nu) \tag{1}$$

oscillator strength $f_{\ell u}$ usually known (i.e., measured in lab)
so at high resolution:

- transmission profile *depth* $\rightarrow$ absorber *column density* N_ℓ
- transmission profile *shape* $\rightarrow$ absorber profile function $\phi_{\ell u}(\nu)$ which encodes, e.g., temperature via core width $b = \sqrt{2kT/m}$, and collisional broadening via wing with Γ

Depth of Line Center

if the absorbers have a Gaussian velocity distribution
then the optical depth profile is $\tau_\nu = \tau_0 \; e^{-v^2/b^2}$
with the Doppler velocity $v = (\nu_0 - \nu)/\nu_0 \; c$, and thus
τ_ν is also Gaussian in ν

the optical depth a the line center is

$$\tau_0 = \sqrt{\pi} \left(\frac{e^2}{m_e c} \right) \; \frac{N_\ell f_{\ell u} \lambda_{\ell u}}{b} \; \left[1 - \frac{g_u N_u}{g_\ell N_\ell} \right] \tag{2}$$

ignoring the stimulated emission term $[\cdots]$, for H Lyman α

$$\tau_0 = 0.7580 \; \left(\frac{N_\ell}{10^{13} \text{ cm}^{-2}} \right) \; \left(\frac{f_{\ell u}}{0.4164} \right) \left(\frac{\lambda_{\ell u}}{1215.7 \text{ Å}} \right) \left(\frac{10 \text{ km/s}}{b} \right)$$

so if we can measure τ_0, we get column N_ℓ

Q: in low-resolution spectra, what sets transmission profile?
Q: what information is lost? what information remains?

Equivalent Width

if instrumental resolution $R = \Delta\lambda_{\rm inst}/\lambda$ low: $\Delta\lambda_{\rm inst} \ll$ line shape
$\rightarrow$ all information about true astrophysical line profile is lost!
and observed profile is just instrumental artifact

yet flux is still removed by the absorption line
so that we still can measure *integrated* effect of line
i.e., the total flux "lost" due to absorbers
$\Delta F_{\rm line} = \int_{\Delta\nu_{\rm line}} [F_\nu(0) - F_\nu] \, d\nu$
where ν_0 is frequency of *line center*

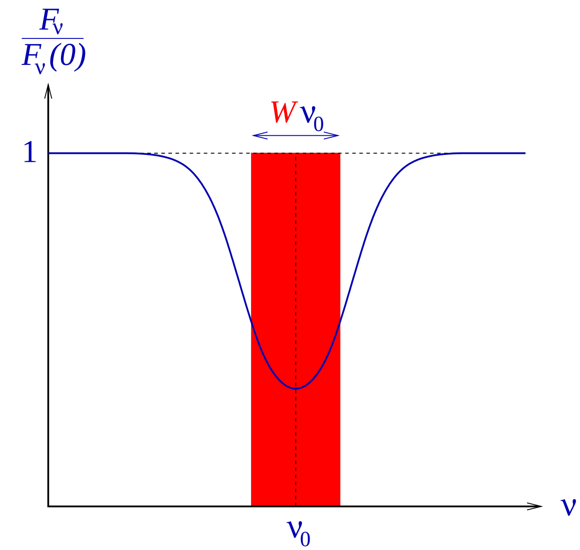
useful to define a dimensionless **equivalent width**

$$W \equiv \frac{\Delta F_{\rm line}}{\nu_0 \; F_\nu(0)} = \int_{\Delta\nu_{\rm line}} \frac{F_\nu(0) - F_\nu}{F_\nu(0)} \; \frac{d\nu}{\nu_0} \tag{3}$$

Q: what does this correspond to physically?

equivalent width

$$W = \int_{\Delta\nu_{\rm line}} \frac{F_\nu(0) - F_\nu}{F_\nu(0)} \, \frac{d\nu}{\nu_0}$$



so $W\nu_0$ equivalent to

width of 100% absorbed line

i.e., *saturated* line with “rectangular” profile

and W is width as fraction of ν_0

note: many authors use *dimensionful equivalent* with

$$W \equiv \frac{W_\lambda}{\lambda_0} = \int_{\Delta\lambda_{\rm line}} \frac{F_\lambda(0) - F_\lambda}{F_\lambda(0)} \, \frac{d\lambda}{\lambda_0} \qquad (4)$$

so that $W_\lambda \approx \Delta\lambda \approx \lambda_0 W$

or the *velocity equivalent width* $W_v = c\, W$

Curve of Growth

in terms of optical depth, equivalent width is

$$W = \int_{\Delta\nu_{\rm line}} \left[1 - \frac{F_\nu}{F_\nu(0)}\right] \frac{d\nu}{\nu_0} = \int_{\Delta\nu_{\rm line}} \left(1 - e^{-\tau_\nu}\right) \frac{d\nu}{\nu_0} \qquad (5)$$

and thus $W = W(N_\ell)$ via $\tau_\nu = \sigma_\nu N_\ell$
dependence of W vs N_ℓ: **curve of growth**

even if line is unresolved, equivalent width still measures
$\Delta F = W \nu_0 F_\nu(0) =$ *total missing flux* across the line

Q: what is W if absorbers are optically thin? what do we learn?

Optically Thin Absorption: $\tau_0 \lesssim 1$

for an optically thin line: $\tau_0 \lesssim 1$
and thus maximal flux reduction at line center is $e^{-\tau_0} \gtrsim 1/e$

if $\tau_\nu \ll 1$ then we can put $1 - e^{-\tau_\nu} \approx \tau_\nu$:

$$W \approx \int \tau_\nu \frac{d\nu}{\nu_0} = N_\ell \frac{\int_{\text{line}} \sigma_{\ell u}(\nu)\ d\nu}{\nu_0} \tag{6}$$

so $W \propto N_\ell$: *linear regime* in curve of growth

for Gaussian profile, good fit to second order in τ_0 is

$$W \approx \sqrt{\pi}\ \frac{b}{c} \frac{\tau_0}{1 + \tau_0/(2\sqrt{2})} = \frac{\pi e^2}{m_e c^2}\ N_\ell\ f_{\ell u}\ \lambda_{\ell u} \frac{\tau_0}{1 + \tau_0/(2\sqrt{2})} \tag{7}$$

and thus when $\tau_0 \ll 1$,

$$N_\ell = \frac{m_e c^2}{\pi e^2} \frac{W}{f_{\ell u}\ \lambda_{\ell u}} = 1.130 \times 10^{12}\ \text{cm}^{-1}\ \frac{W}{f_{\ell u}\ \lambda_{\ell u}} \tag{8}$$

if line optically thin, then $W \propto N_\ell$
width measures absorber column density

sketch W *vs* N_ℓ

Q: what happens if line is optically thick?

Q: what if line is thick and we assume thin?

Q: how can we use W *to check if line is thick or thin?*

Flat Part of Curve of Growth: $1 \lesssim \tau_0 \lesssim \tau_{\rm damp}$

once $\tau_0 \gtrsim 1$, line center has essentially no flux
$\rightarrow$ line *core* is totally dark and thus *saturated*
true line profile is nearly "*box-shaped*"

true line shape still has damping wings
but there cross section is small, so if $\tau_0 \lesssim \tau_{\rm damp}$
then wings only "round the edges" of the line "box"

if we treat the *unresolved* line as a box
then width is just Gaussian width

$$W \approx \frac{(\Delta\nu)_{\rm FWHM}}{\nu_0} = \frac{(\Delta v)_{\rm FWHM}}{c} = \frac{2\ b}{c}\sqrt{\frac{\ln \tau_0}{2}} \tag{9}$$

and thus $W \propto b\ \sqrt{\ln \tau_0}$
Q: implications?

when $1 \lesssim \tau_0 \lesssim \tau_{\rm damp}$ then equivalent width
$W \propto b\sqrt{\ln \tau_0}$ depends very weakly on N_ℓ
$\rightarrow$ "flat part" of curve of growth
add to W vs N_ℓ sketch

solving for column:

$$N_\ell \approx \frac{\ln 2}{\sqrt{\pi}} \frac{m_e c}{e^2} \frac{b}{f_{\ell u} \lambda_{\ell u}} e^{(cW/2b)^2} \tag{10}$$

column is *exponentially sensitive* to W

Warning! if a line is in this regime:

- difficult to get N_ℓ from measurements of W
- reliable result requires
  - ▷ very accurate measurements of W *and* b
  - ▷ confidence that true line profile is Gaussian

Q: what if absorber column density increases further?

Damped Part of Curve of Growth: $\tau_0 > \tau_{\rm damp}$

if N_ℓ and thus τ_0 very large,
then absorption very strong, then high-res profile
shows *Lorentzian "damping wings"*

away from line center, in "wing" regime $|\nu - \nu_0| \gg \nu_0/b/c$:

$$\tau_\nu \approx \frac{\pi e^2}{m_e c} N_\ell f_{\ell u} \; \frac{4\Gamma_{\ell u}}{16\pi^2(\nu-\nu_0)^2 + \Gamma_{\ell u}^2} \qquad (11)$$

full width at half-max, i.e., width at 50% transmission, is

$$\frac{(\Delta\lambda)_{\rm FWHM}}{\lambda_0} = \frac{(\Delta u)_{\rm FWHM}}{c} = \sqrt{\frac{1}{\pi \ln 2} \frac{e^2}{m_e c} N_\ell f_{\ell u} \lambda_{\ell u} \frac{\Gamma_{\ell u}}{\nu_{\ell u}}}$$

thus equivalent width is

$$W = \sqrt{\pi \ \ln 2} \frac{(\Delta\lambda)_{\text{FWHM}}}{\lambda_0} = \sqrt{\frac{e^2}{m_e c} N_\ell f_{\ell u} \lambda_{\ell u} \frac{\Gamma_{\ell u}}{\nu_{\ell u}}} = \sqrt{\frac{b}{c} \frac{\tau_0}{\sqrt{\pi}} \frac{\Gamma_{\ell u} \lambda_{\ell u}}{c}} \tag{12}$$

finish W *vs* N_ℓ *sketch*

`www: professional plot of curve of growth`

so the column density is

$$N_\ell = \frac{m_e c^3}{e^2} \frac{W^2}{f_{\ell u} \Gamma_{\ell u} \lambda_{\ell u}^2} \tag{13}$$

transition from flat to damped when $W_{\text{flat}} \approx W_{\text{dampled}}$:

$$\tau_{\text{damp}} \approx 4\sqrt{\pi} \ \frac{b}{\Gamma_{\ell u} \lambda_{\ell u}} \ln \left[\frac{4\sqrt{\pi}}{\ln 2} \frac{b}{\Gamma_{\ell u} \lambda_{\ell u}} \right] \tag{14}$$

Awesome Example: Quasar Absorption Lines

Q: let's remind ourselves–what's a quasar?

quasar (QSO) rest-frame optical to UV spectra $F_\lambda(0) = F_\lambda^{\mathrm{qso}}$:
• *smooth continuum*, with
• *broad peak* at rest-frame *Lyman-α* line
`www: high-resolution quasar spectrum`

quasars generally at large redshift, *typically* $z_{\mathrm{qso}} \sim 3$
• distance very large: $\gtrsim d_H \sim 4000$ Mpc
• observed peak at $\lambda_{\mathrm{peak,obs}} = (1+z_{\mathrm{qso}})\lambda_{\mathrm{Ly}\alpha} \sim 3600$ Å: *optical!*
QSO light passes through all intervening material at $z < z_{\mathrm{qso}}$

Q: what is intervening material made of?
Q: effect if absorbers have uniform comoving cosmic density?
Q: why can we rule out a uniform density?

Quasar Absorption Line Systems

quasars are distant, high-redshift *backlighting*
to all of the foreground universe

but thanks to big-bang nucleosynthesis, we know:
cosmic *baryonic** matter mostly made of *hydrogen*

if universe *uniformly filled* with H in $1s$ ground state, then:

- *at redshift* z, Lyα $1s \to 2s$ absorption
  at absorber-frame $\lambda_{\mathrm{Ly}\alpha}$, and observer-frame $\lambda_{\mathrm{obs}} = (1+z)\lambda_{\mathrm{Ly}\alpha}$
  absorption should occur at all $\lambda < (1 + z_{\mathrm{qso}})\lambda_{\mathrm{Ly}\alpha}$
- absorbers have same comoving density at each z
  so optical depth τ_λ and hence transmission *spectrum*
  should be *smooth* as a function of λ

*in cosmo-practice: a *baryon* = *neutron* or *proton* or combinations of them
= *anything made of atoms* = *ordinary matter* ≠ dark matter

Observed quasar spectra:

- *do* show absorption shortwards of the quasar Lyα!
- but transmitted spectrum is not smooth continuum, rather, a series of many separate *lines*

Implications:

- diffuse intervening neutral hydrogen exists!
  $\rightarrow$ there is an **intergalactic medium**
- intergalactic neutral gas is not uniform but *clumped* into "clouds" of atomic hydrogen

note: low-z quasars show few absorption lines
high-z quasars show many: **Lyman-α forest**
a major cosmological probe

Q: what information does each forest line encode?

The Lyman-α Forest: Observables

each forest *line* $\leftrightarrow$ *cloud of neutral hydrogen*

- absorber z_{abs} gives *cloud redshift*
- absorber depth gives cloud *column density* N(H I)

note that absorbers span wide range in column densities

- most common: optically thin "*forest systems*"
  correspond to *modest overdensities* $\delta\rho/\rho \sim 1$
- rare: optically thick "*damped Lyα systems*"
  damping wings of seen in line profile $\rightarrow N(\mathrm{H\ I}) \gtrsim 10^{20}\ \mathrm{cm}^{-2}$
  correspond to *large overdensities: protogalaxies!*

`www: zoom into damped Ly`α `system`